

Core-Collapse Supernova Explosions and Their Gravitational Wave Signatures

*Adam Burrows, Viktoriya Morozova,
David Vartanyan, David Radice,
Aaron Skinner, Josh Dolence*

Supported by:
DOE/SciDAC4
NSF/MPPC
NSF/AST

FORNAX: 1D,2D,3D, Multi-Group,
Radiation/Hydrodynamics

FORNAX: 1D,2D,3D, Multi-Group, Explicit Radiation/Hydrodynamics

- ◆ Solves the Two-Moment Transport Equations, with 2nd and 3rd moment closures (not “ray-by-ray”); **second-order accurate in space and time**
- ◆ Explicit **Riemann Godunov-like** solution to the Transport operator
- ◆ Terms of $O(v/c)$ included in transport
- ◆ **Implicit** solution to the local transport source terms
- ◆ **Explicit** hydro; full energy and momentum couplings – HLLC
- ◆ **Conserves** energy, momentum, and lepton number to machine precision
- ◆ **Very good** energy conservation with gravity included
- ◆ “6” – Dim. = 1(time) + 3(space) + 1(energy-group) + vector Flux
- ◆ Logically spherical coordinates – **general metric/covariant formulation**
- ◆ **Multipole Gravity** (can include GR-like modifications to the monopole)
- ◆ Multi-D calculated to the center - Core refinement (“dendritic grid”) – improves timestepping by many factors (!); **static mesh refinement**
- ◆ For 2D, Axisymmetry – **Rotation can be included (conserving angular momentum to machine precision)**
- ◆ Good **strong scaling** in core count and scaling in energy group (linear)
- ◆ **Result: Faster multi-D supernova code (by factor of ~5-10)**

FORNAX (cont.)

- ◆ Includes: **Inelastic scattering** off **electrons**
- ◆ **Inelastic scattering** off **nucleons**
- ◆ Includes in-medium **Many-body response** corrections (Horowitz et al. 2017)
- ◆ **General-relativistic** monopole gravity correction and **gravitational redshifts** (can compare with Newtonian)
- ◆ Multi-D transport, with rbr+ option (for comparison)
- ◆ Weak magnetism and recoil corrections
- ◆ Multipole gravity (with monopole variant)

The Fornax Supernova Code

Dynamical Equations

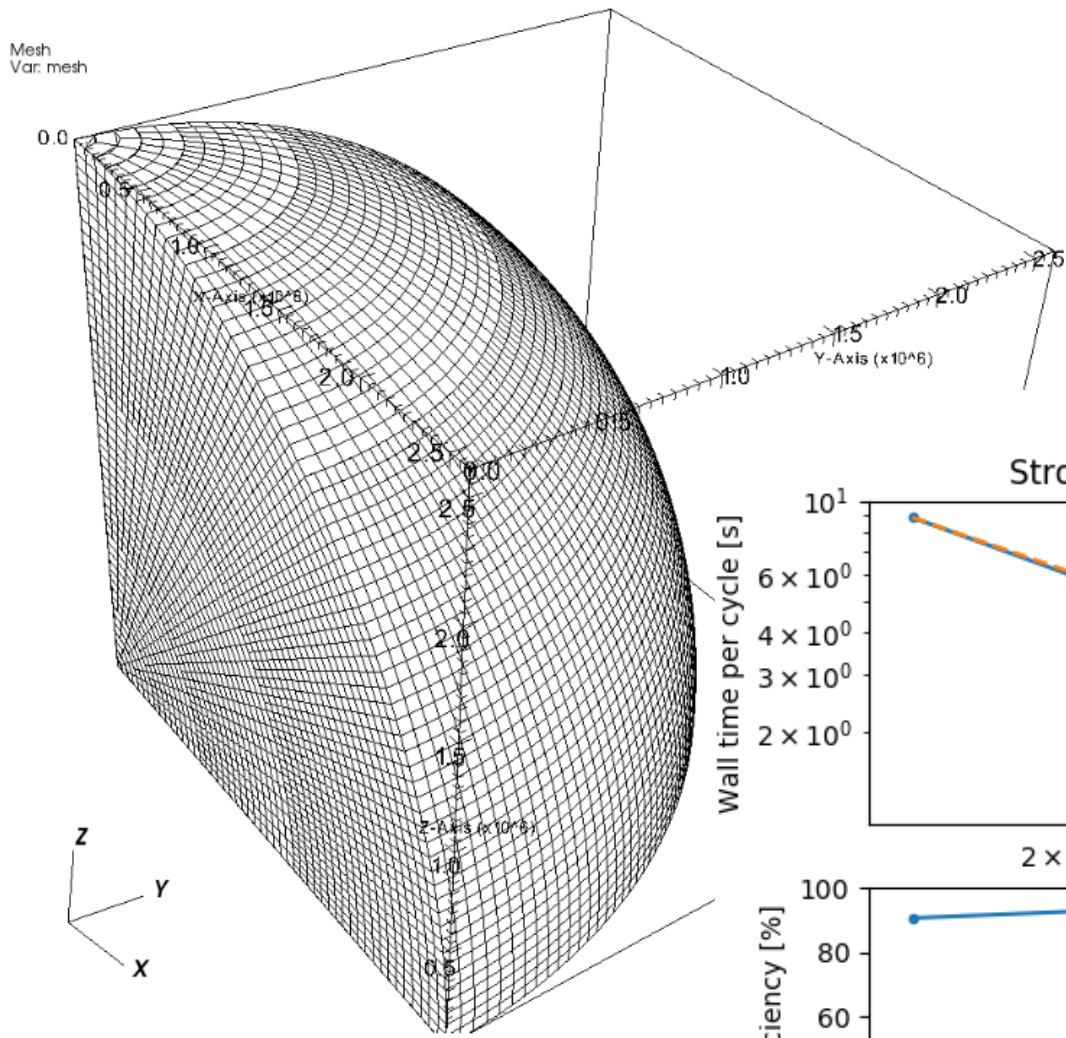
$$\begin{aligned}
 \rho_{,t} + (\rho v^i)_{;i} &= 0 \\
 (\rho v_j)_{,t} + (\rho v^i v_j + P \delta_j^i)_{;i} &= -\rho \phi_{,j} + c^{-1} \sum_s \int_0^\infty (\kappa_{s\varepsilon} + \sigma_{s\varepsilon}^{\text{tr}}) F_{s\varepsilon j} d\varepsilon \\
 \left[\rho \left(e + \frac{1}{2} \|v\|^2 \right) \right]_{,t} + \left[\rho v^i \left(e + \frac{1}{2} \|v\|^2 + \frac{P}{\rho} \right) \right]_{;i} &= -\rho v^i \phi_{,i} \\
 &\quad - \sum_s \int_0^\infty \left(j_{s\varepsilon} - c \kappa_{s\varepsilon} E_{s\varepsilon} - \frac{v^i}{c} (\kappa_{s\varepsilon} + \sigma_{s\varepsilon}^{\text{tr}}) F_{s\varepsilon i} \right) d\varepsilon \\
 (\rho Y_e)_{,t} + (\rho Y_e v^i)_{;i} &= \sum_s \int_0^\infty \xi_{s\varepsilon} (j_{s\varepsilon} - c \kappa_{s\varepsilon} E_{s\varepsilon}) d\varepsilon \\
 E_{s\varepsilon,t} + (\alpha F_{s\varepsilon}^i + \mathbf{v}^i \mathbf{E}_{s\varepsilon})_{;i} - \alpha \mathbf{v}_{;j}^i \frac{\partial}{\partial \ln \varepsilon} \mathbf{P}_{s\varepsilon i}^j &= \alpha (j_{s\varepsilon} - c \kappa_{s\varepsilon} E_{s\varepsilon}) + \alpha G^e \\
 F_{s\varepsilon j,t} + (c^2 \alpha P_{s\varepsilon j}^i + \mathbf{v}^i \mathbf{F}_{s\varepsilon j})_{;i} + \alpha \mathbf{v}_{;j}^i \mathbf{F}_{s\varepsilon i} - \alpha \mathbf{v}_{;k}^i \frac{\partial}{\partial \varepsilon} (\varepsilon \mathbf{Q}_{s\varepsilon j i}^k) &= -\alpha (\kappa_{s\varepsilon} + \sigma_{s\varepsilon}^{\text{tr}}) F_{s\varepsilon j} + \alpha G_j^m
 \end{aligned}$$

Approximate GR

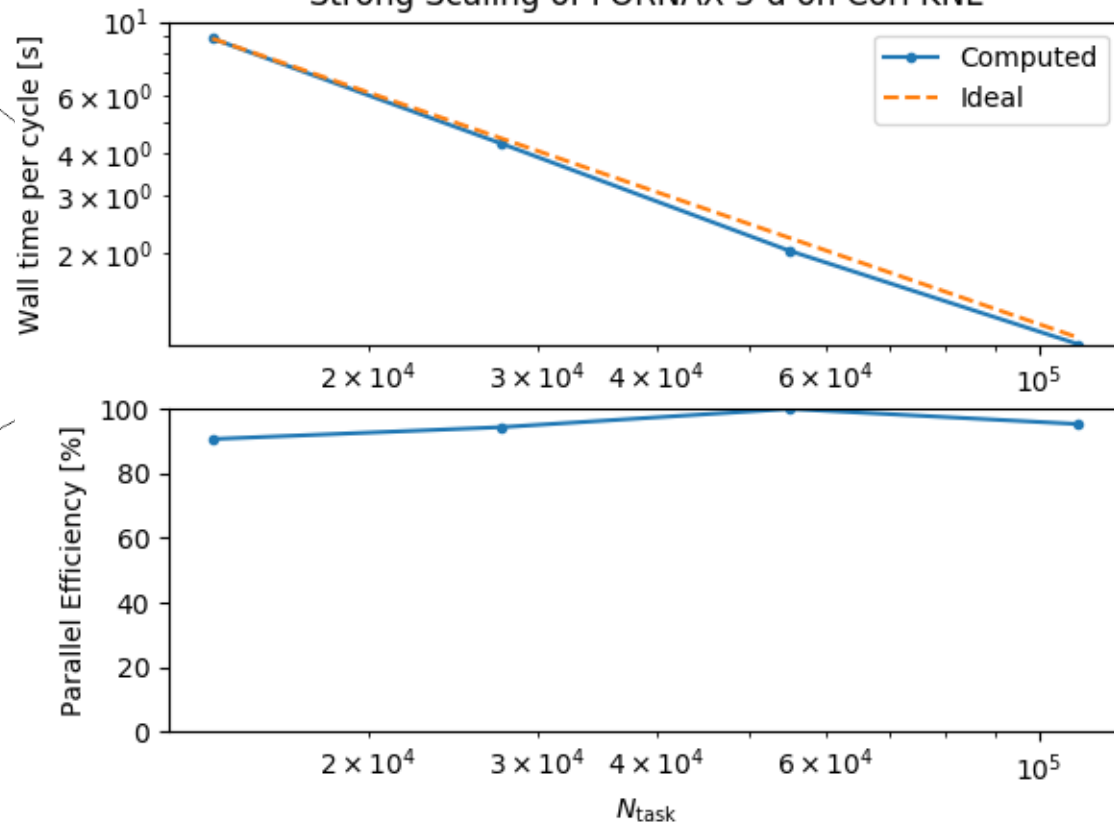
$$\begin{aligned}
 \frac{d\phi_{\text{TOV}}}{dr} &= G \frac{m_{\text{TOV}} + 4\pi r^3 (P + P_\nu)/c^2}{r^2 \Gamma^2} \left[\frac{\rho + E/c^2 + P/c^2}{\rho} \right] \\
 \Gamma(r) &= \sqrt{1 + v^2/c^2 - \frac{2Gm_{\text{TOV}}(r)}{rc^2}} \\
 \frac{dm_{\text{TOV}}}{dr} &= 4\pi r^2 \left(\rho + E/c^2 + E_\nu/c^2 + \frac{\mathbf{v} \cdot \mathbf{F}_\nu/c^2}{\Gamma} \right) \Gamma
 \end{aligned}$$

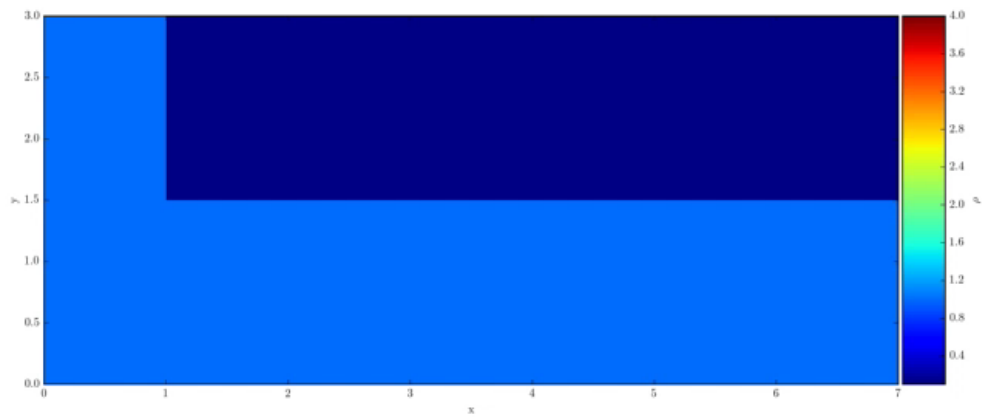
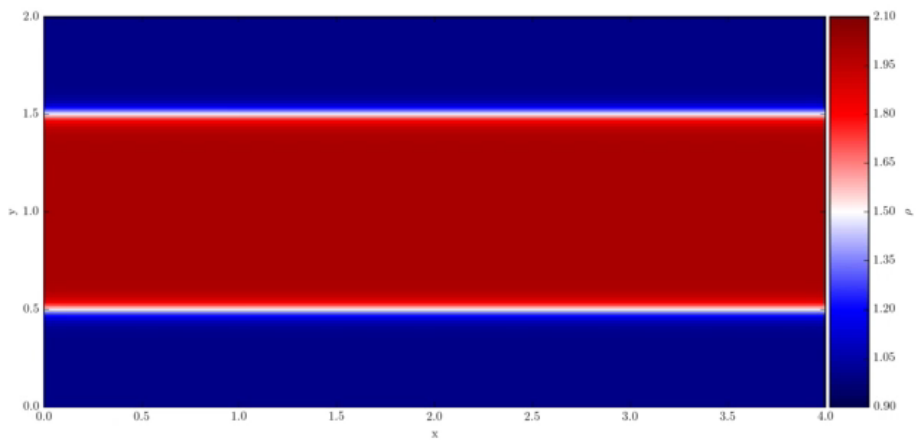
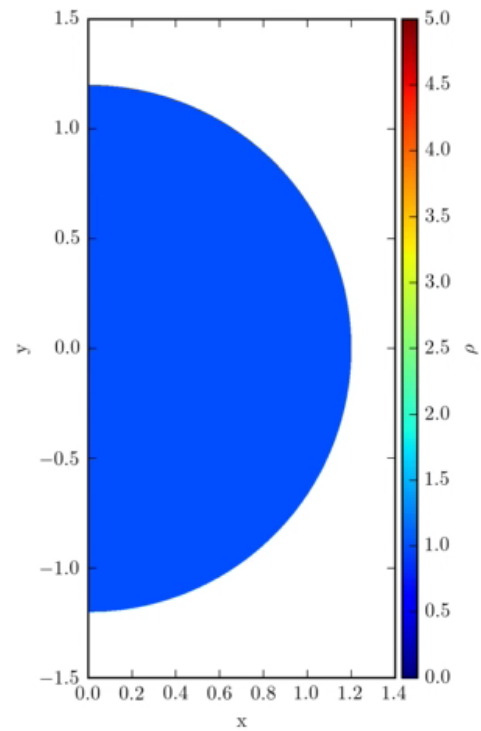
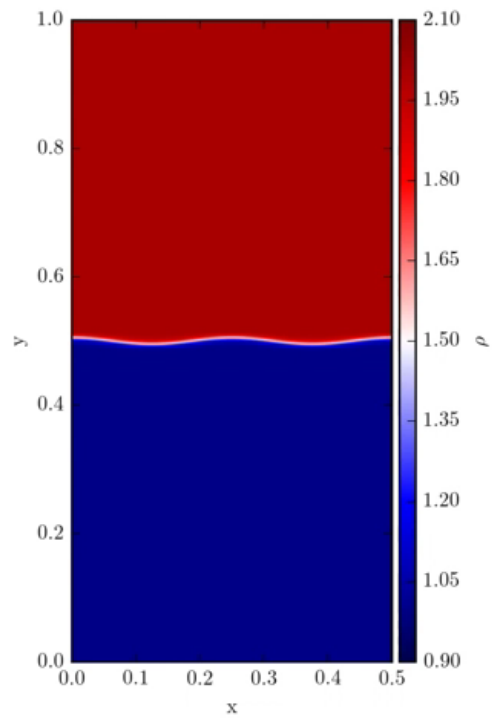
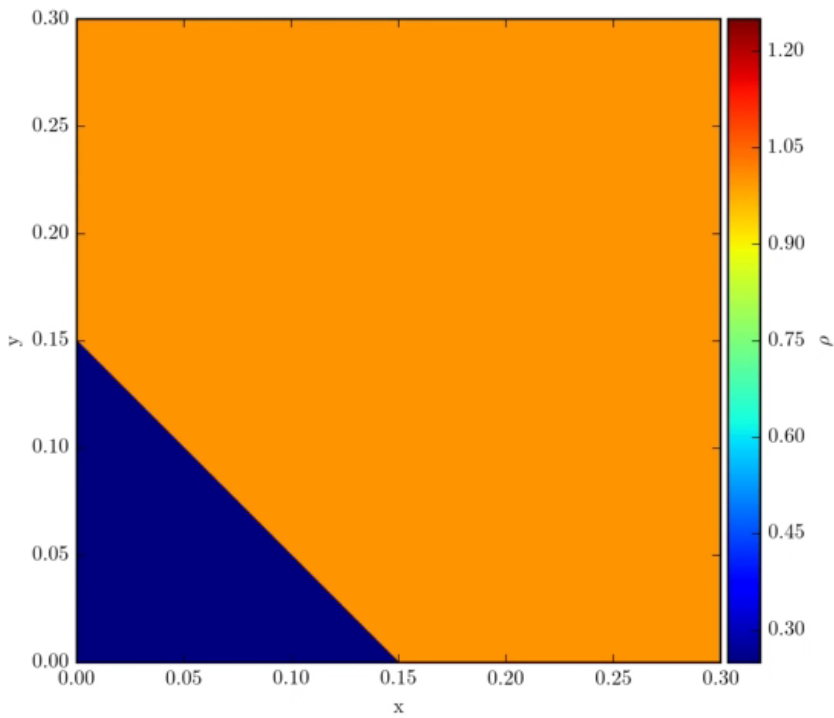
Auxiliary Definitions

$$\begin{aligned}
 \alpha &= e^{\phi/c^2} \\
 G^e &= -\mathbf{F}_{s\varepsilon} \cdot \nabla \phi/c^2 + \nabla \phi/c^2 \cdot \partial(\varepsilon \mathbf{F}_{s\varepsilon})/\partial \varepsilon \\
 G_j^m &= -E_{s\varepsilon} \nabla_j \phi/c^2 + \nabla_i \phi/c^2 \cdot \partial(\varepsilon \mathbf{P}_{s\varepsilon j}^i)/\partial \varepsilon \\
 \xi_{s\varepsilon} &= \begin{cases} -(N_A \varepsilon)^{-1} & s = \nu_e, \\ (N_A \varepsilon)^{-1} & s = \bar{\nu}_e, \\ 0 & s = \nu_x \end{cases}
 \end{aligned}$$



Strong Scaling of FORNAX 3-d on Cori KNL





Different Groups get Different Explosion
Times, Energies, Even Outcomes, etc.

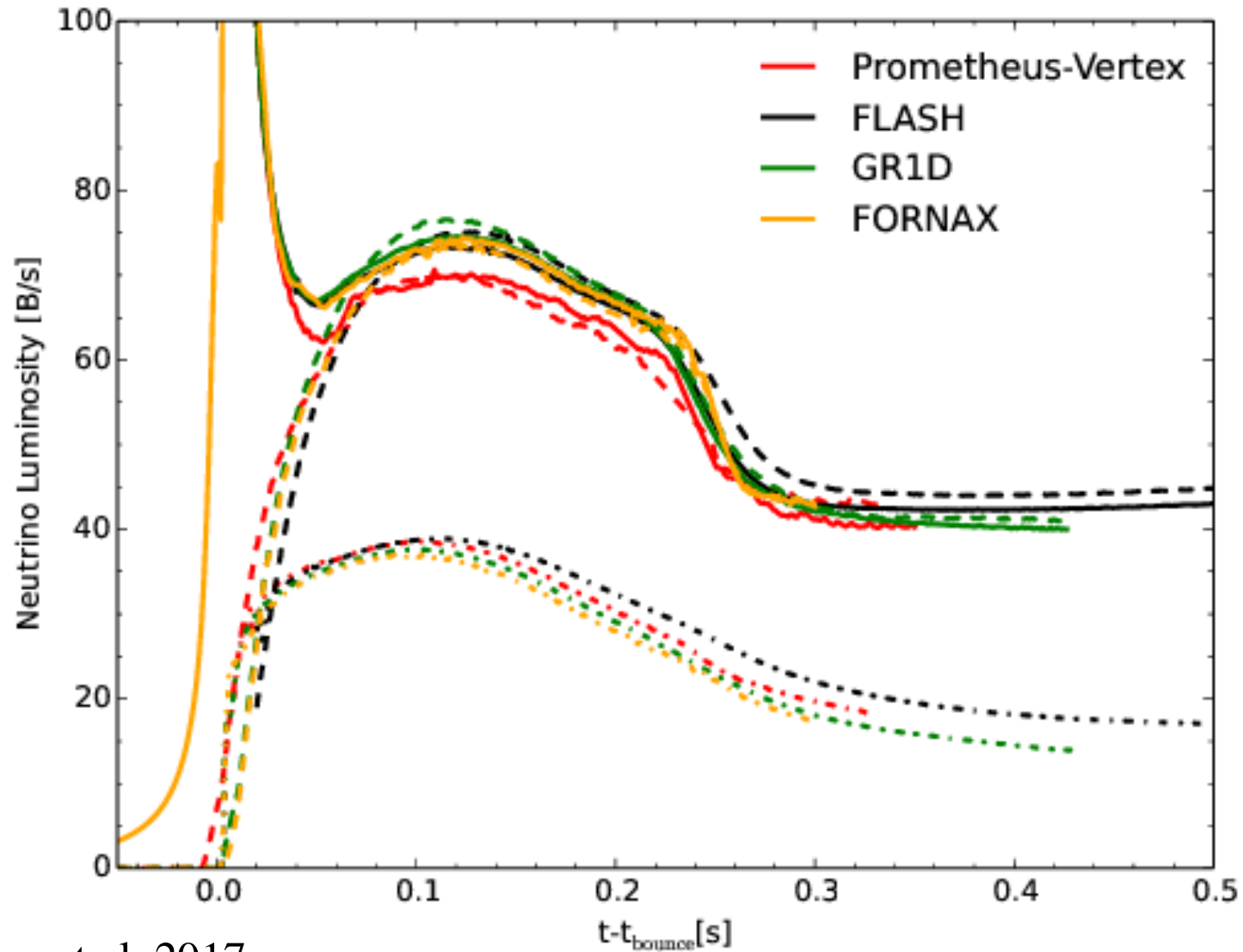
Why?

What are the Crucial Ingredients for a
Robust Explosion?

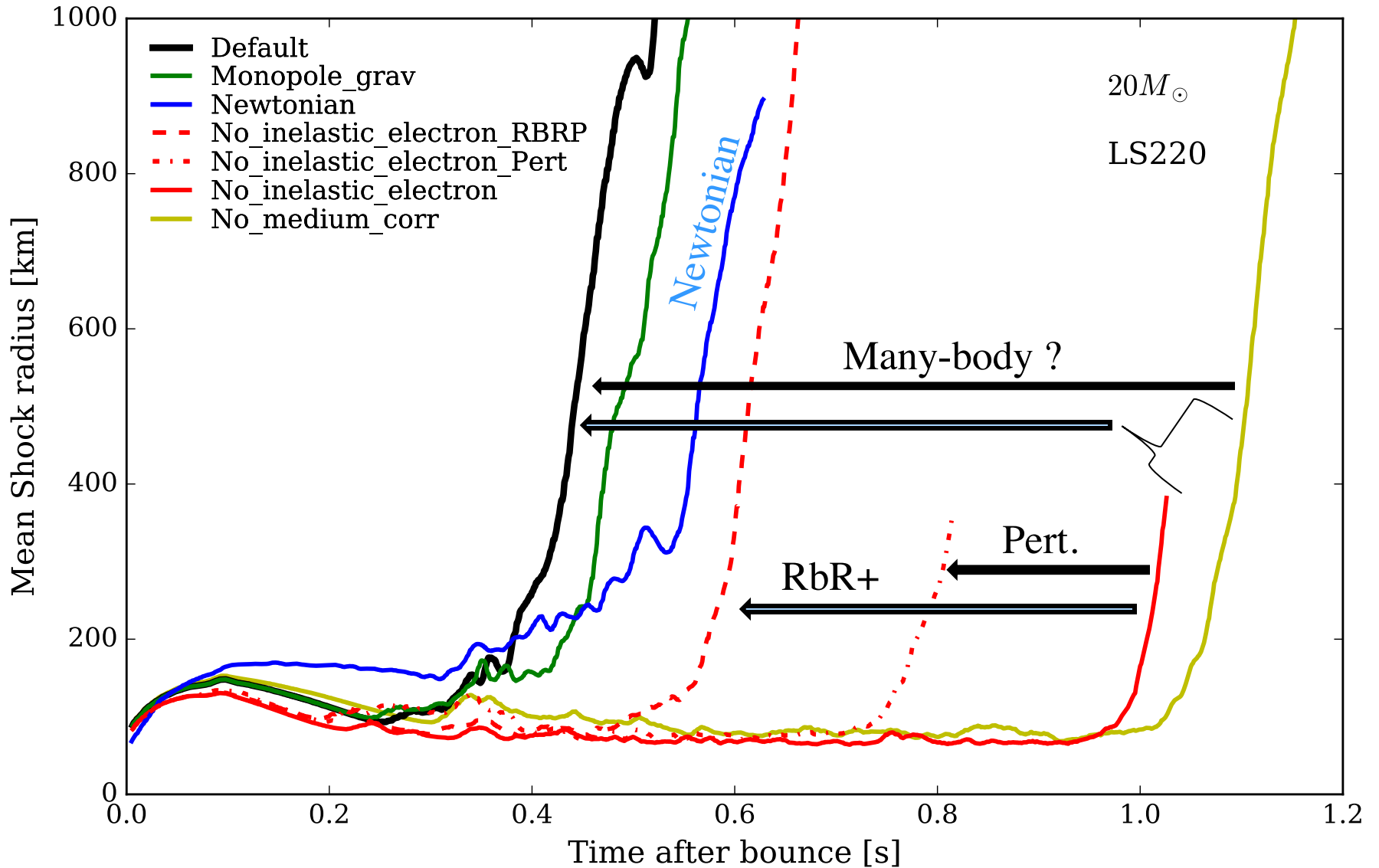
Why are different groups witnessing different outcomes?

- Multi-D simulations are close to critical condition
- Flow is chaotic (positive Lyapunov exponents!)
- Whereas in 1D, there are severe feedbacks to changes in microphysics (“Mazurek’s Law), in multi-D proximity to critical condition amplifies sensitivities to slight changes in interactions, methodology, resolution, etc.
- In-medium (many-body) structure factor/response corrections to neutrino-nucleon scattering (only 10’s of %) - Burrows & Sawyer 1998; C. Horowitz/T. Fischer
- Spatial (!) and energy resolution
- Ray-by-ray+ ?
- Initial progenitor models
- “Seed” Perturbations
- Equation of state (EOS)
- GR/no-GR (strength of gravity -> core compression; redshifts)
- Weak magnetism and recoil corrections to interaction rates
- Inelastic scattering (neutrino-electron; neutrino-nucleon)
- Code/algorithm differences
- Optimization flags, machine compilers, machine precision

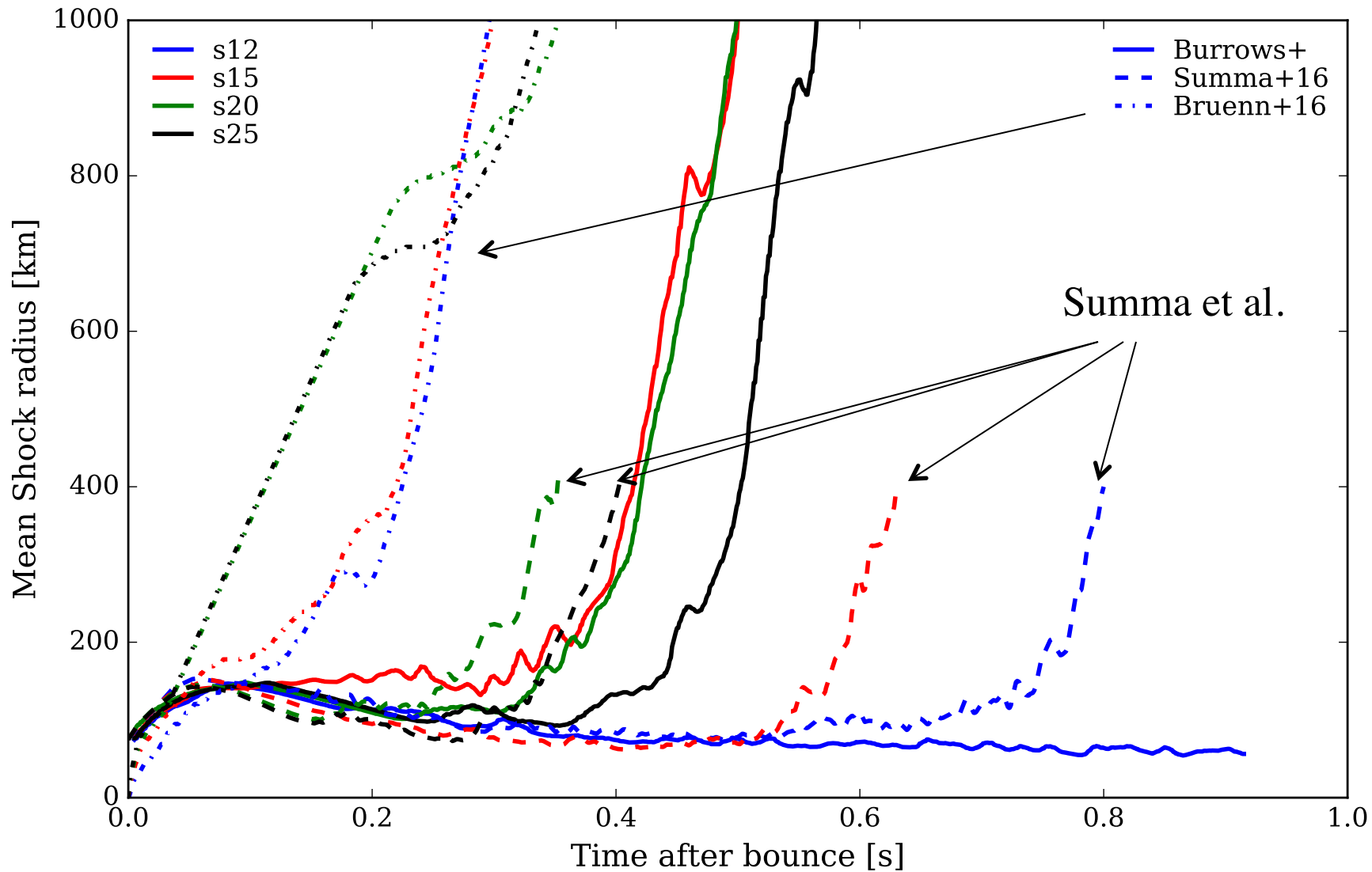
Luminosity Curve Comparisons (1D) – Different Groups;
Sukhbold 2015/20 solar-mass; SFHo EOS, GR, IES



Small Changes and Differences Result in Significantly Different Outcomes



Different Groups Explode at Different Times



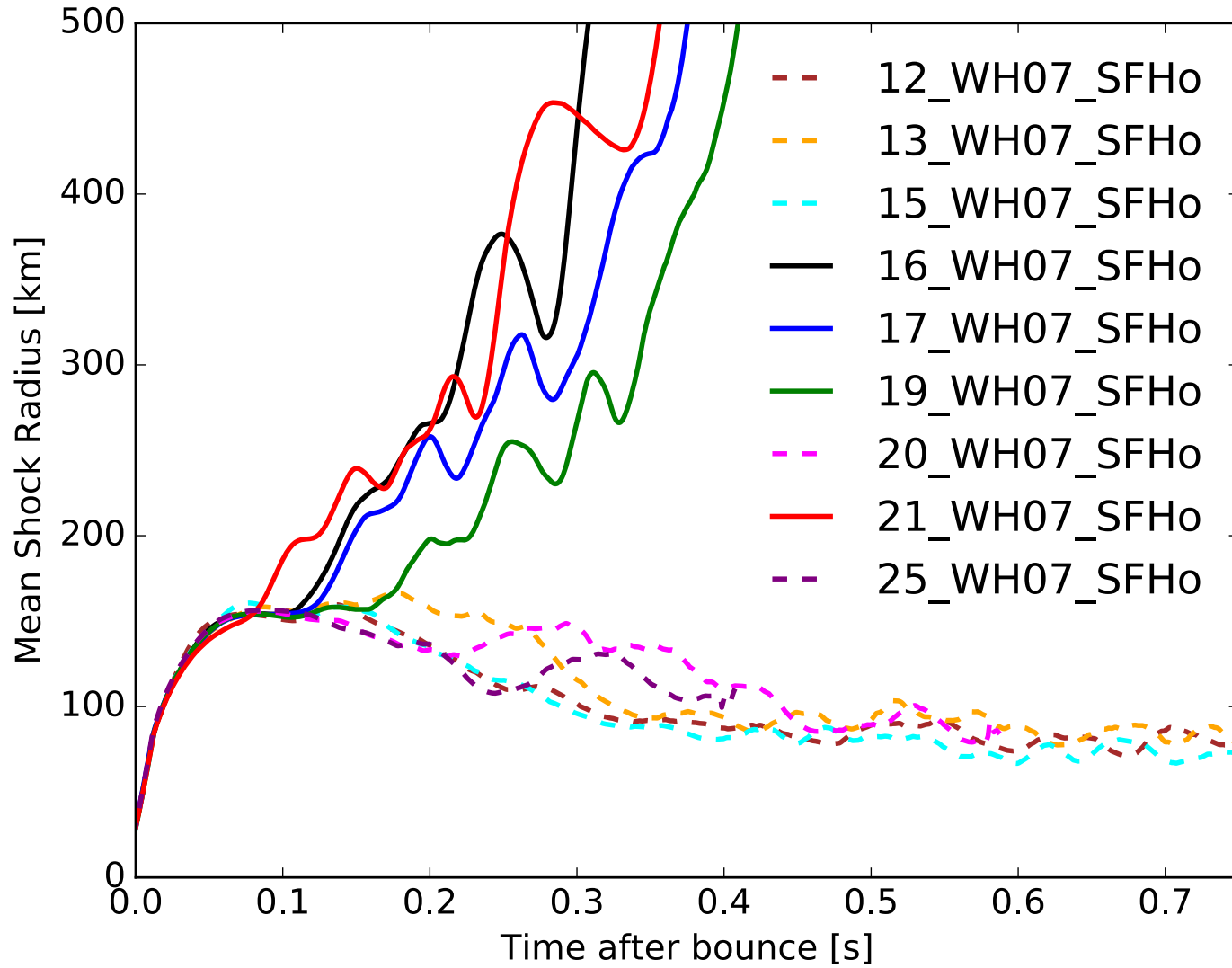
2D (Fornax): 9.0 solar-mass (Sukhbold et al.) model – Without and With Structure Factor/
Many-body Correction – Pressure/Acoustic Waves



Upshot(s):

- Explosions at significantly different post-bounce times (100s of ms!)
- Explosion versus no explosion
- Shock positions chaotic - fluctuate indeterminately
- All issues when inter-comparing in multi-D

Chaos in the turbulent context and **proximity to criticality** lead to a violation of “Mazurek’s Law”: Modest effects can make a difference in explosion time, and explodability



Vartanyan et al. 2017

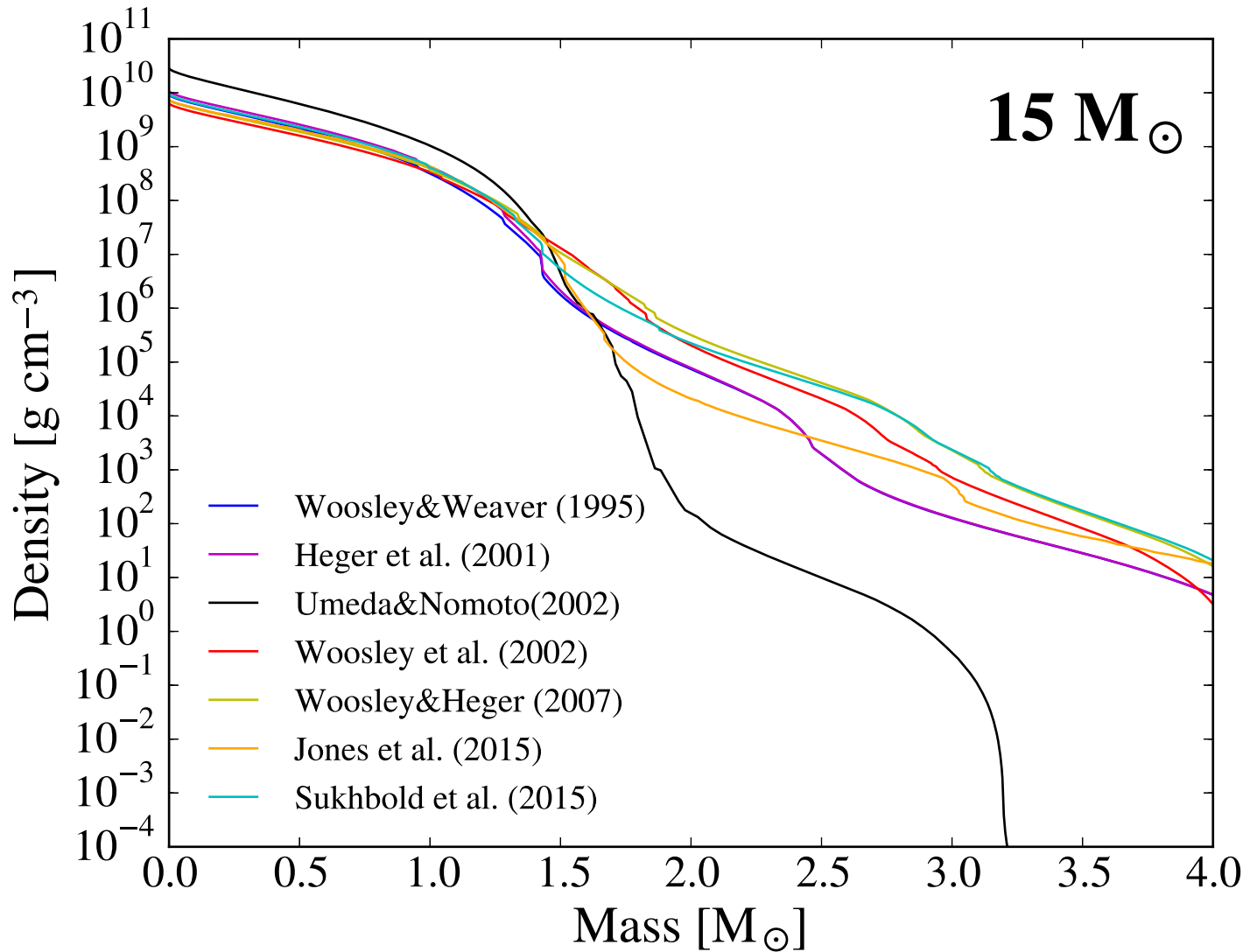
2D (Fornax): 16, 17, 19, 21 solar-mass (Woosley & Heger progenitors) - Baseline

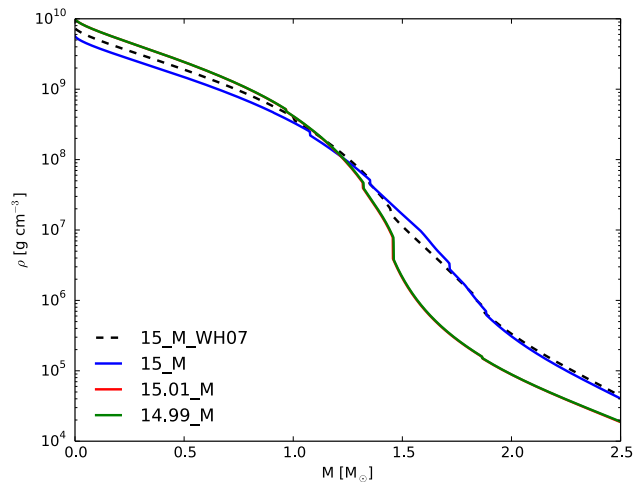
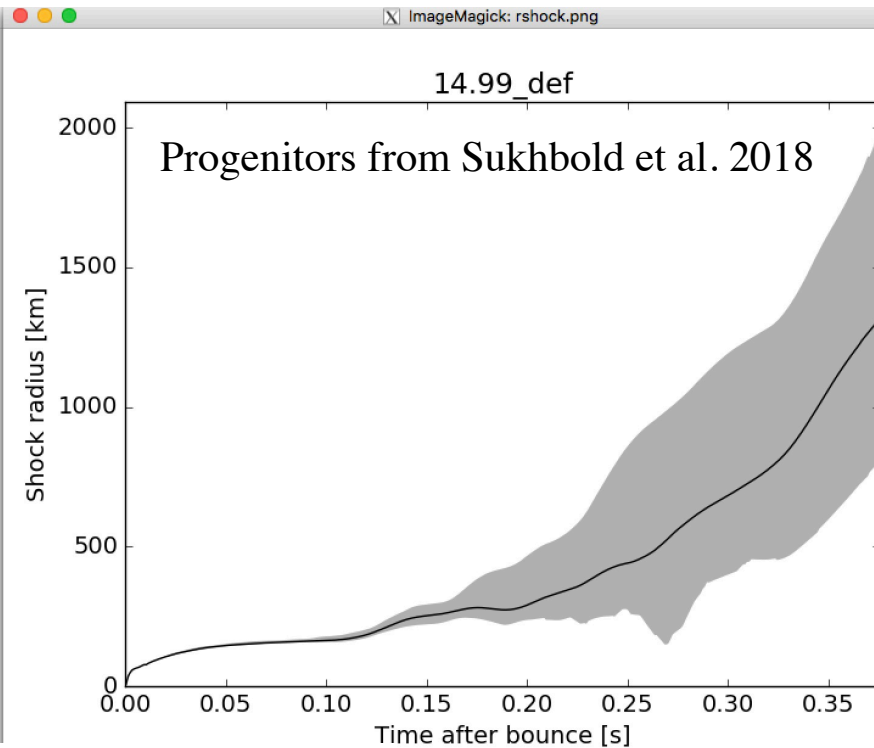
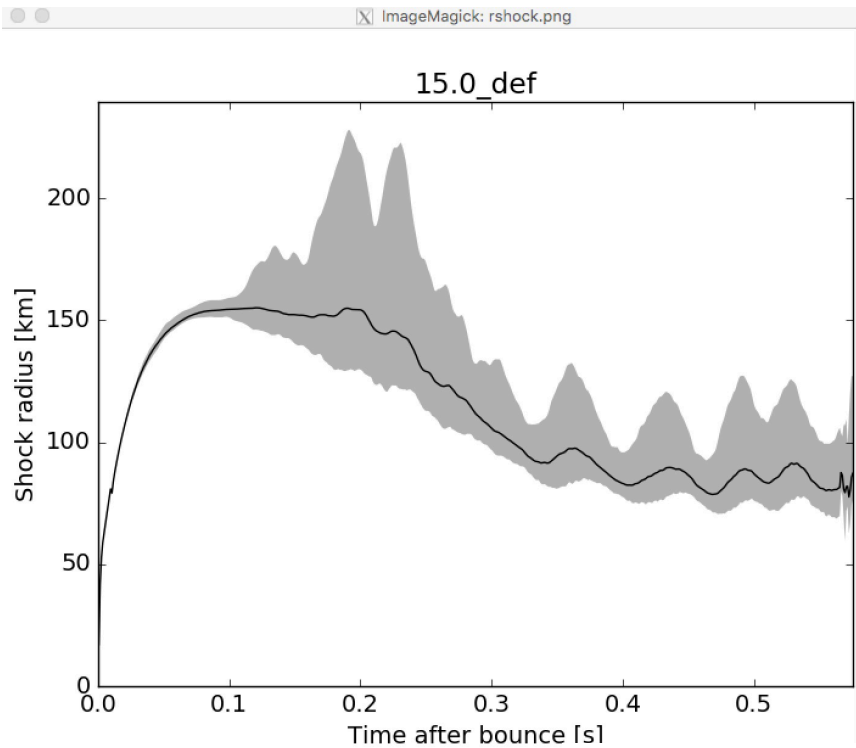


Important Roles of Progenitor Models:

Density Structures, Rotational Profiles,
Seed Perturbations

Different Groups, Same ZAMS Mass





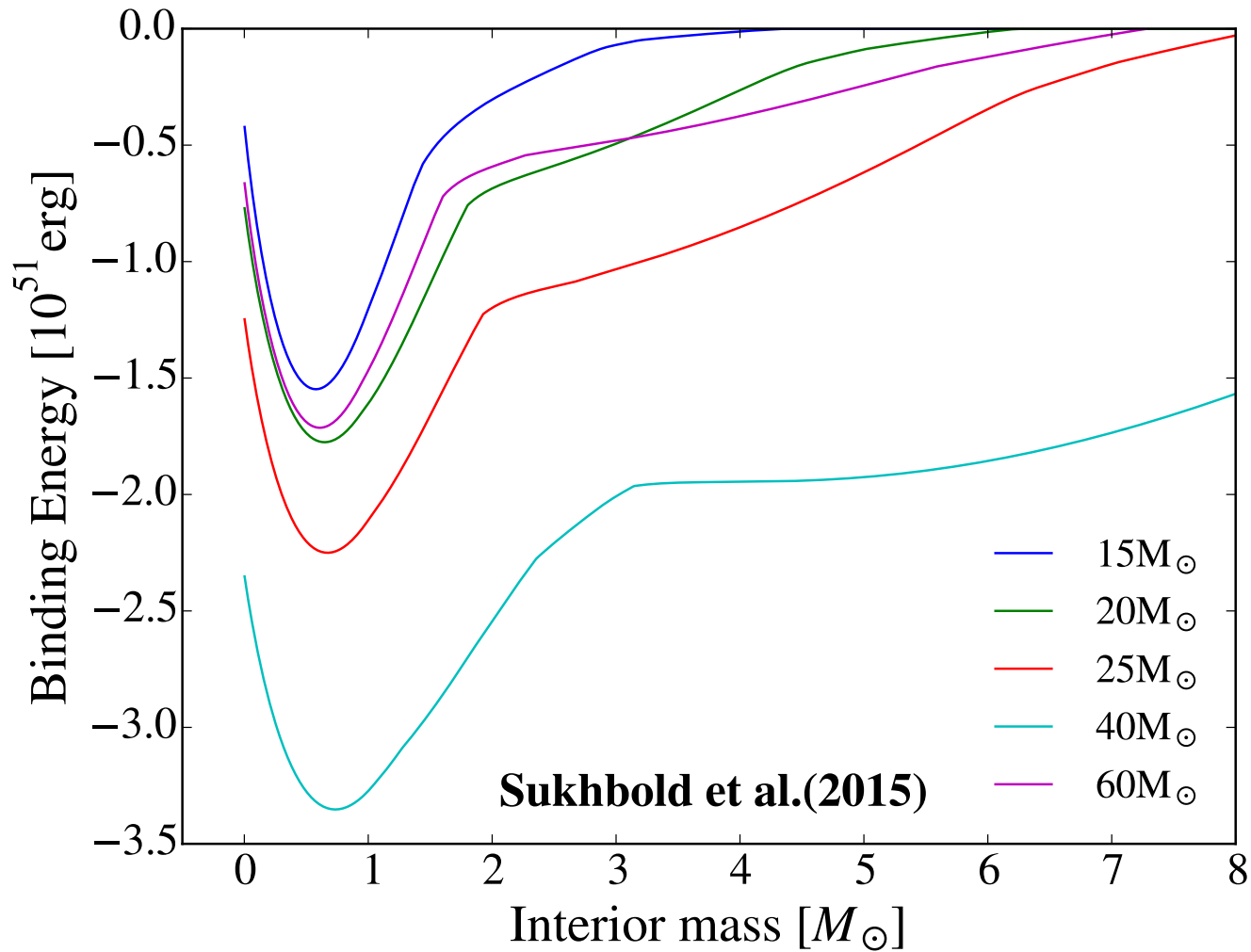
Vartanyan, Burrows, et al. 2018b

Are Progenitor Perturbations/Seeds crucial/
important?

2D (Fornax): Multi-D Transport – 20 solar-mass (WH07) model – With Structure Factor Correction, With and without Perturbations (without inelastic neutrino-electron scattering)



External Binding Energy More Important

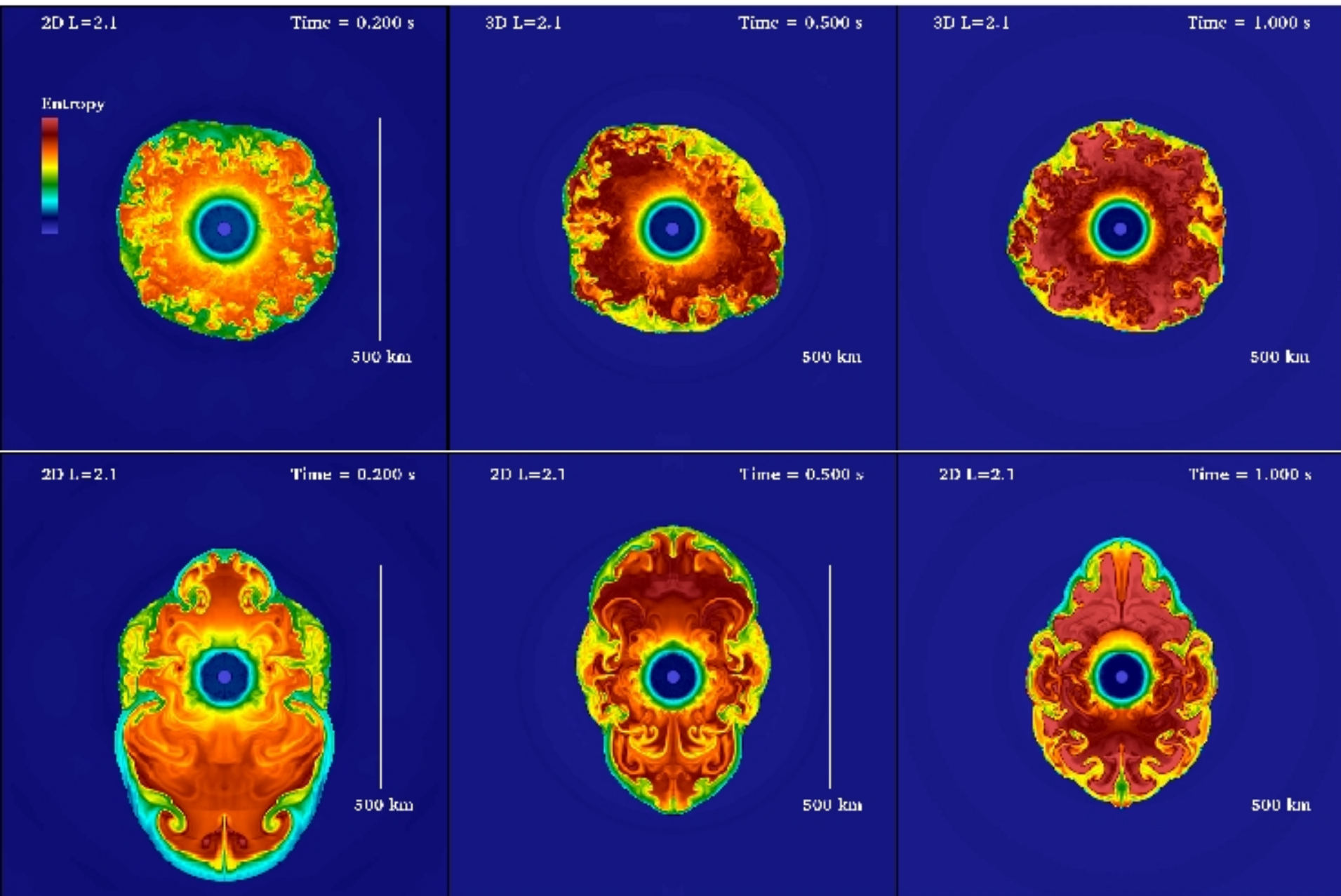


3D Simulations

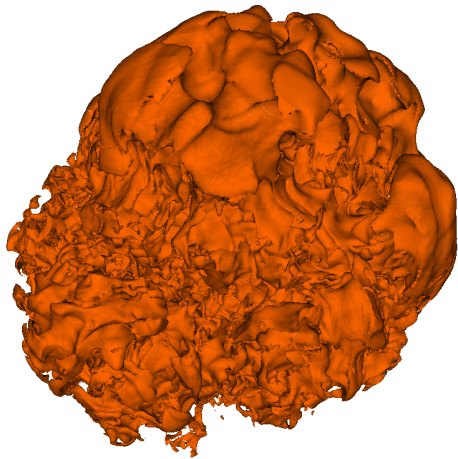
Character of 3D Turbulence
Qualitatively Different than that
in 2D

Amplitudes of Dipolar
("Sloshing") Modes much Smaller
in 3D

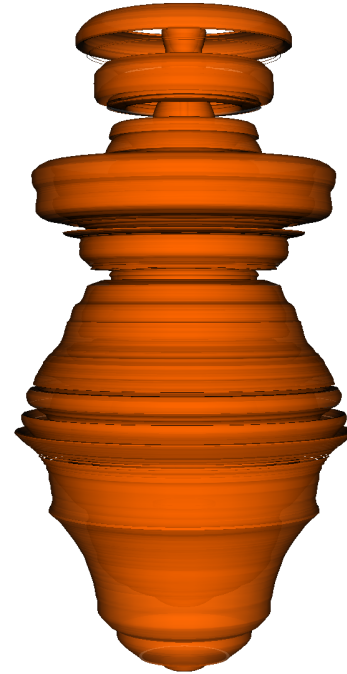
Comparison of 2D with 3D



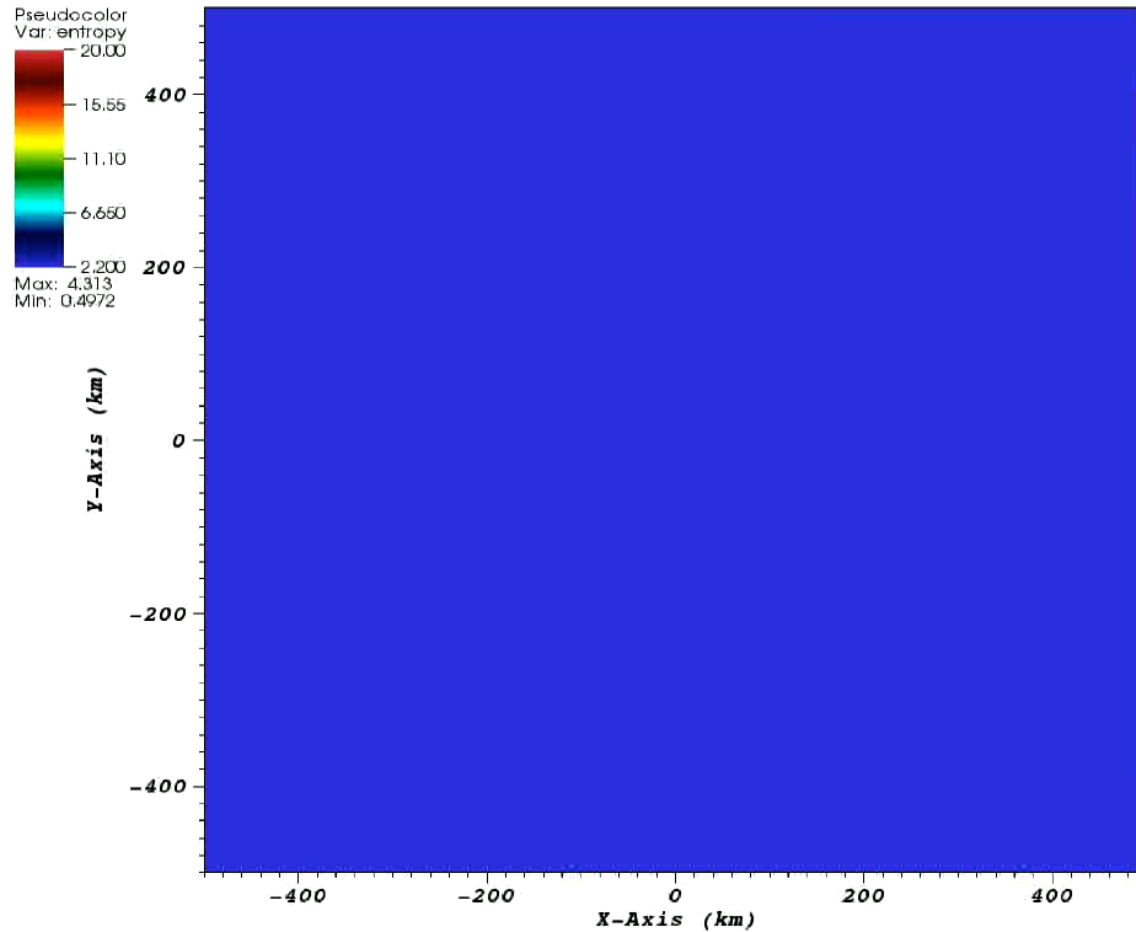
3D



2D



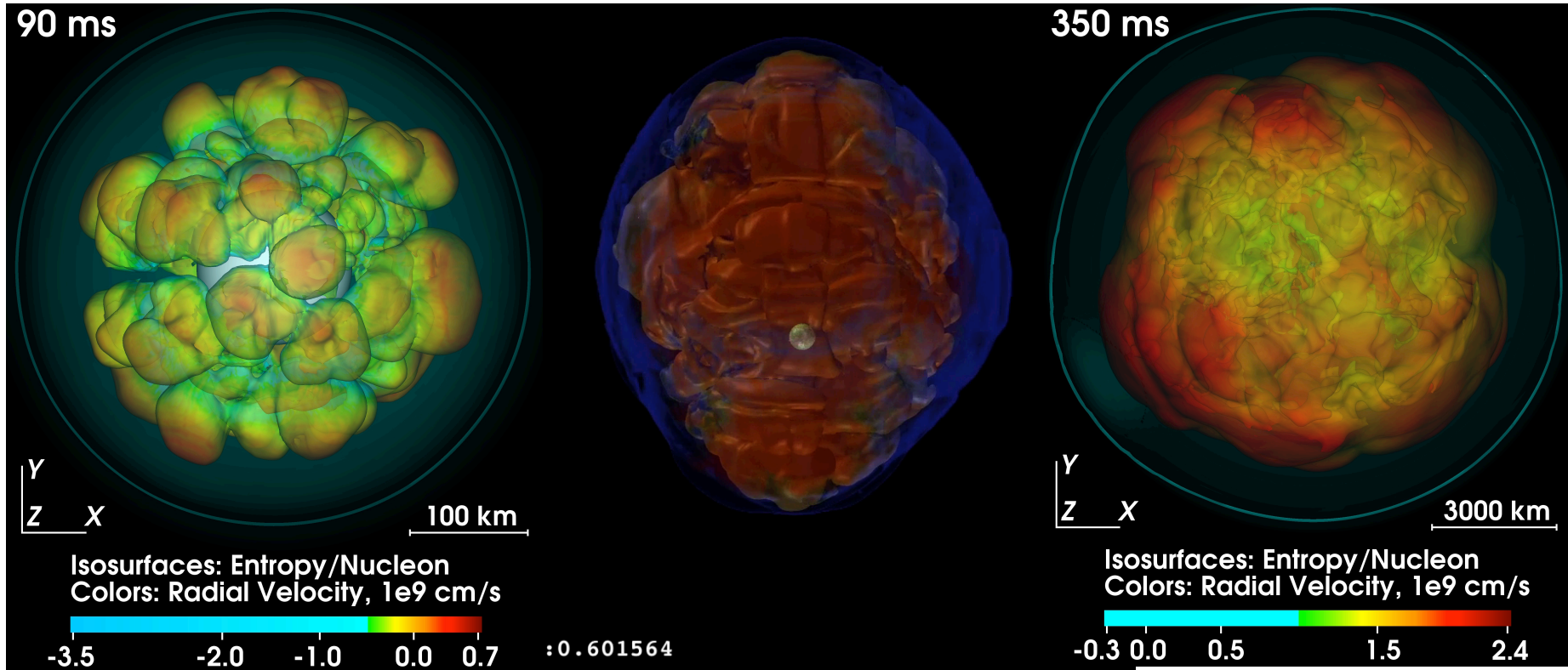
Character of 3D turbulence and Explosion Very Different from those in 2D



Time = -0.256 s after bounce

Melson et al. (2015,
MPA) – 3D (with
strangeness correction -
likely too large)

Lentz et al. (2015) - ORNL



Summa et al. 2017
(MPA) – rapidly
rotating 3D models
(but pulsar birth
spins?)

New Fornax 3D Simulations

Adam Burrows, David Vartanyan, David Radice, Aaron
Skinner, Viktoriya Morozova, Josh Dolence





Gravitational Radiation from Supernovae

Morozova, Radice, Burrows et al. 2018

Identified/Discussed Signal Features

- Rotational **bounce spike** (rapid rotation?); differential rotation
- **Initial** Progenitor perturbation **spike**
- **Outer** PNS convection (early, non-rotating)
- Quiescent **phase** (altered by progenitor perturbations?)
- **Ramp up and saturation of** turbulent convection and SASI
- **Infall** plume excitation of PNS **oscillations**
- Inner PNS convection
- **Transition to Explosion**, leading to decreased accretion, occasioning **signal turnover** (near time of frequency peak?)
- Neutrino **component**
- **Christodoulou Memory** (low frequency): asymmetric explosion, **neutrinos**

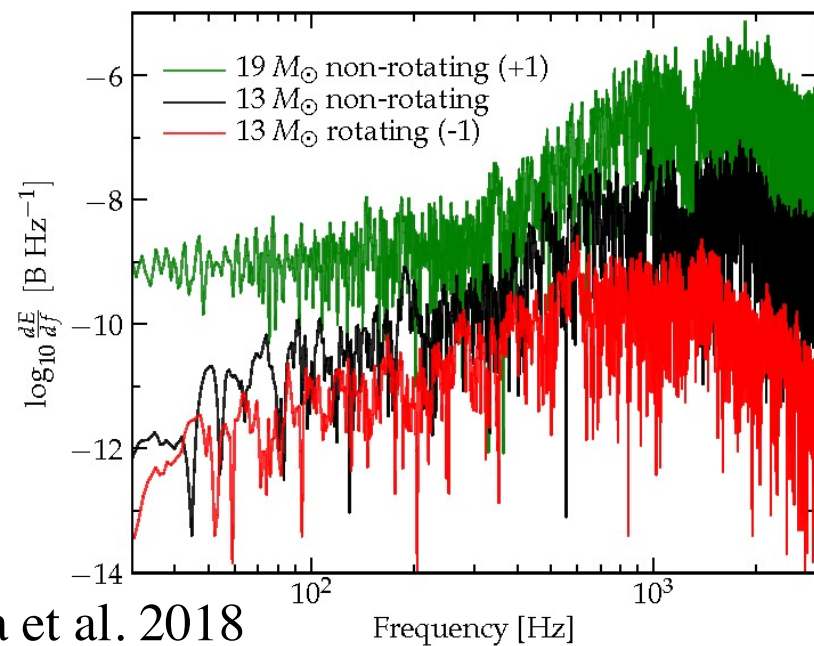
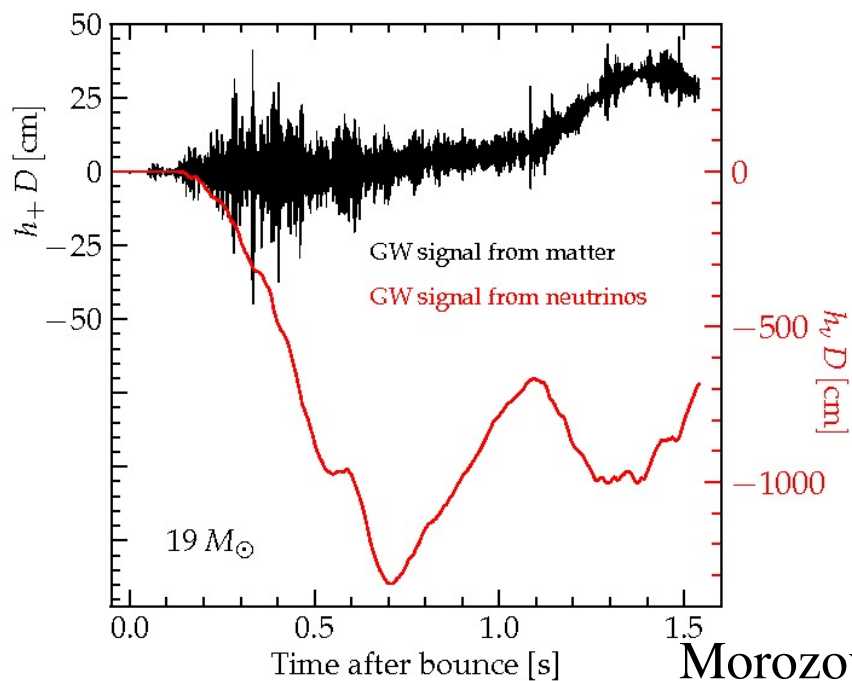
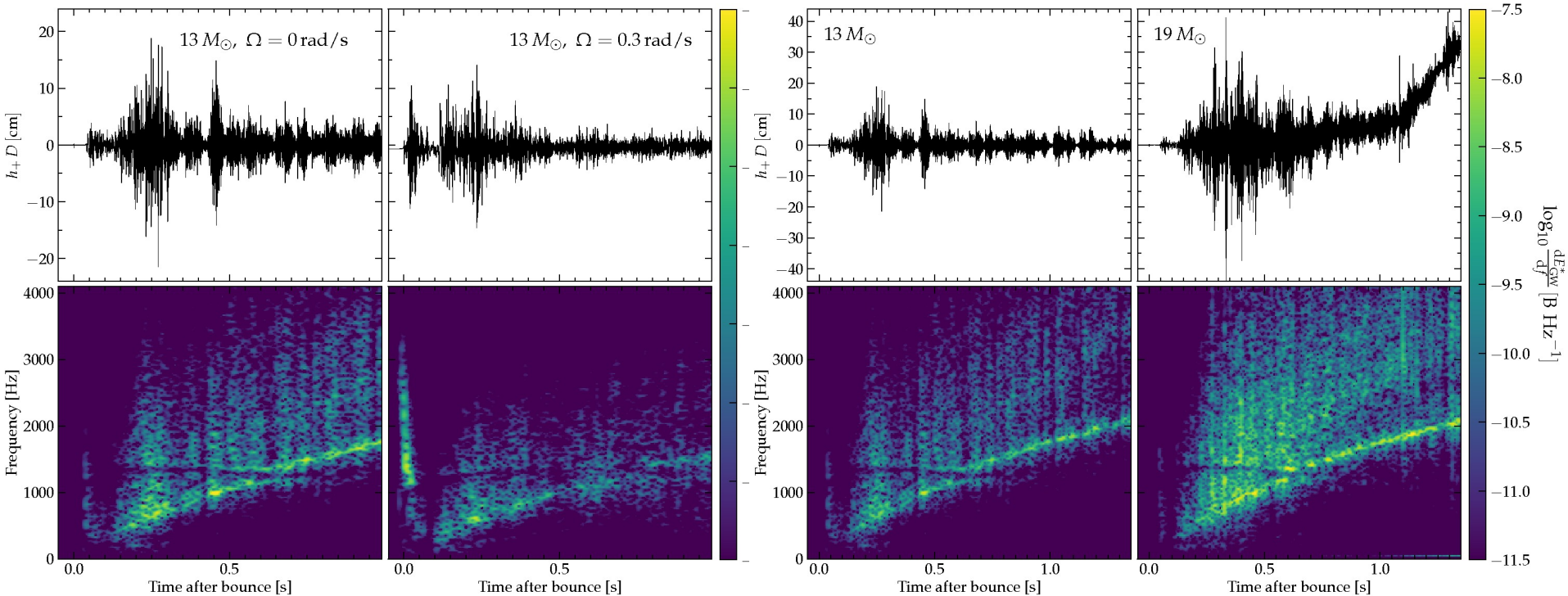
- Progenitor, rotation, orientation, explosion energy dependences?
- Duration of phases; frequency spectra; signal phase?

Sample References on the CCSN/GW Connection

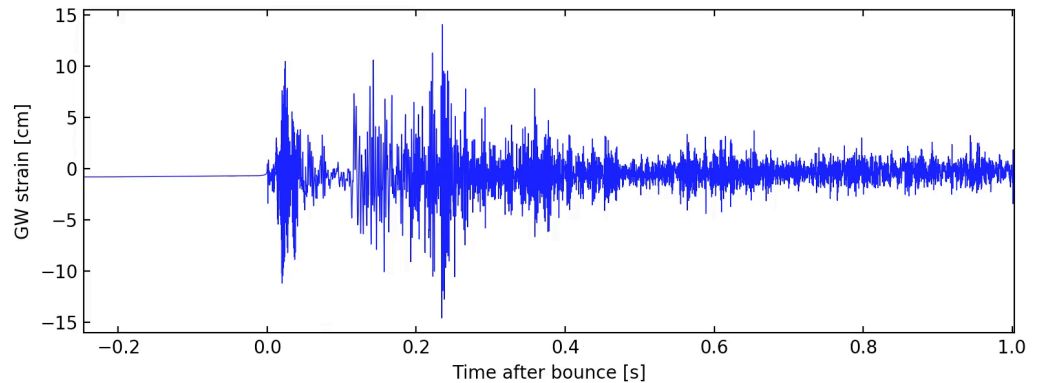
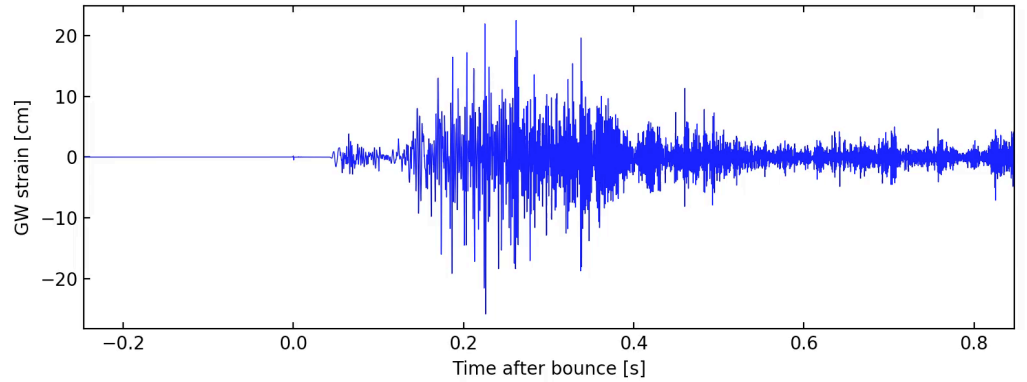
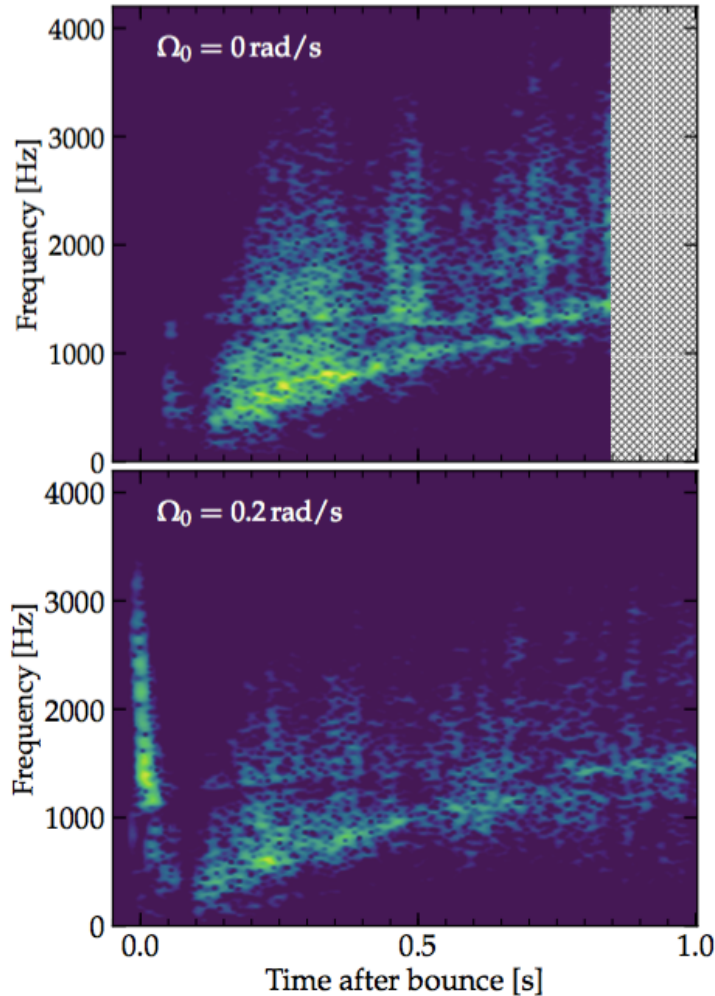
- E. Muller 1982 (Rotation)
- Burrows & Hayes 1996 (Memory, neutrinos, progenitor asymmetry)
- Ott et al. 2004,2006 (Rotation); Dimmelmeier et al. 2008; Ott 2009; Logue et al. 2012; Ott et al. 2013 (3D, leakage)
- Summerscales et al. 2008 (Maximum entropy)
- Scheidegger et al. 2010; Basel
- Heng 2009; Rover et al. 2009
- Murphy et al. 2009 (PNS oscillations, signal sequence)
- Yakunin et al. 2010 (Newt.),2015 (2D),2017(3D); ORNL
- B. Muller et al. 2013 (2D); Andresen et al. 2016 (3D); Garching
- Kuroda et al. 2016 (3D, SASI); Fukuoka
- Gossan et al 2015; Powell et al. 2016 (Signal analysis); Caltech
- Richers et al. 2017 (EOS, rotation); Caltech
- Lynch et al. 2017 (oLIB, Bayes factor); MIT



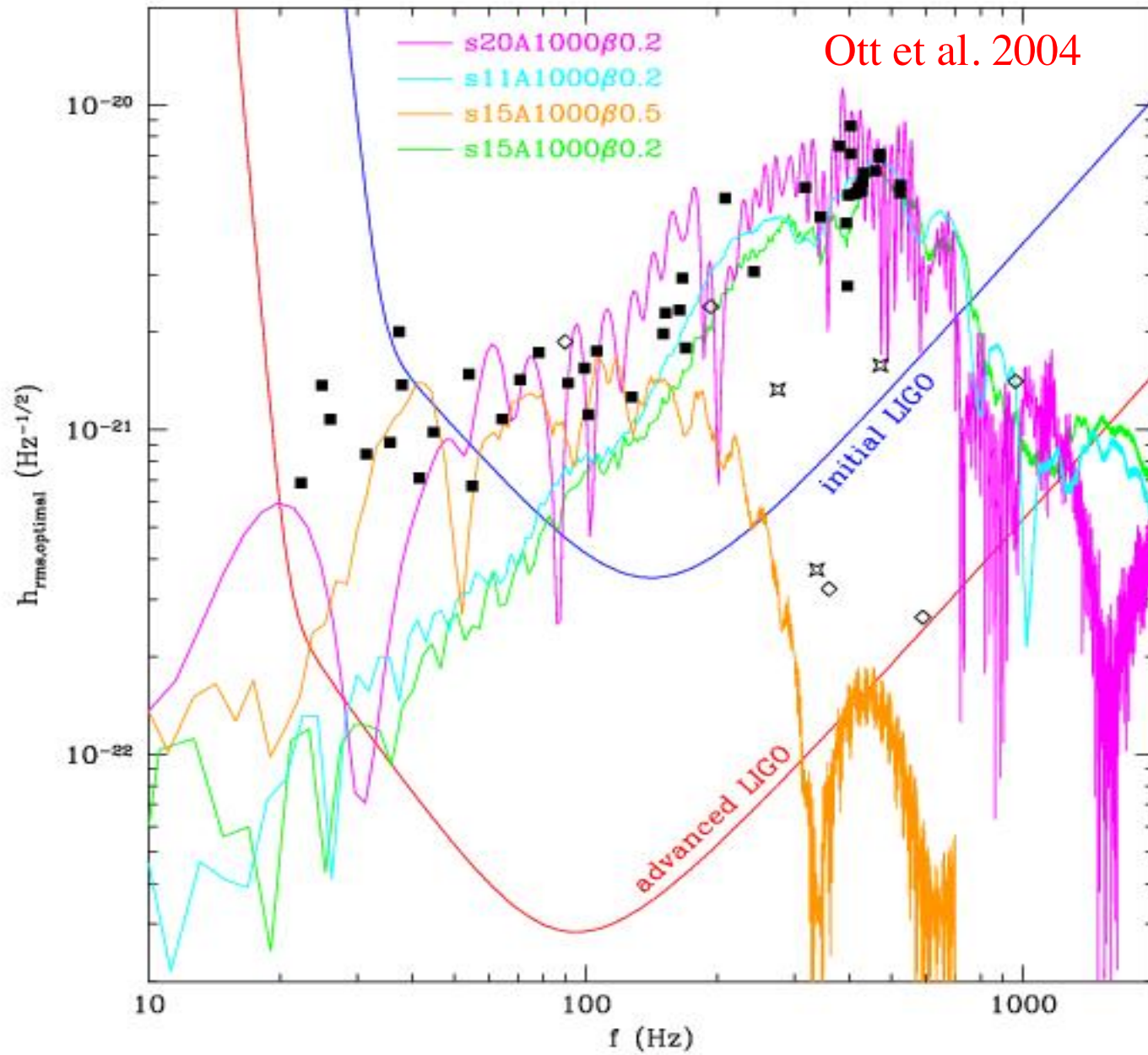




Rotation

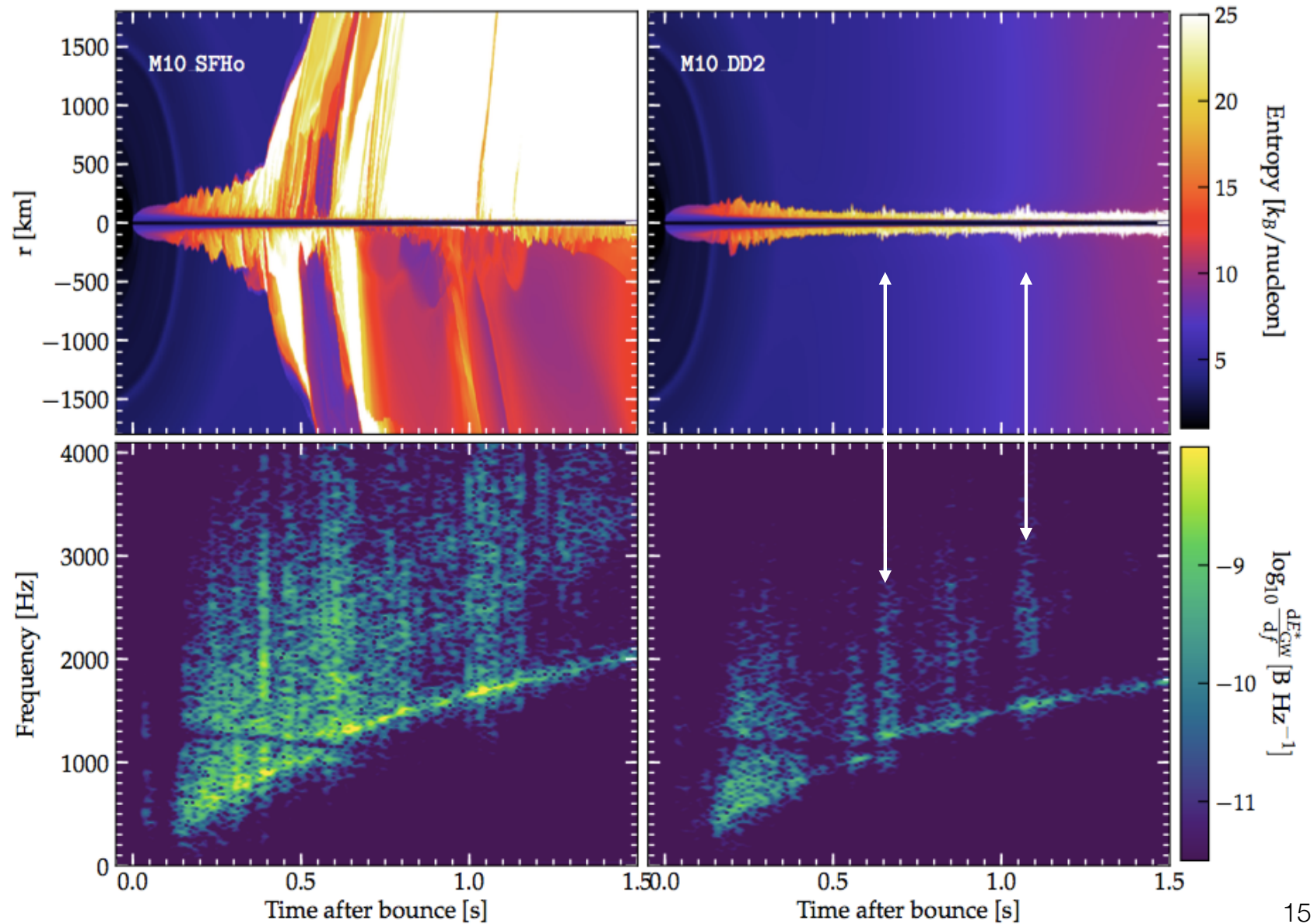


- * The bounce signal is stronger, because the collapse is not symmetric
- * The dominant frequency is nearly the same

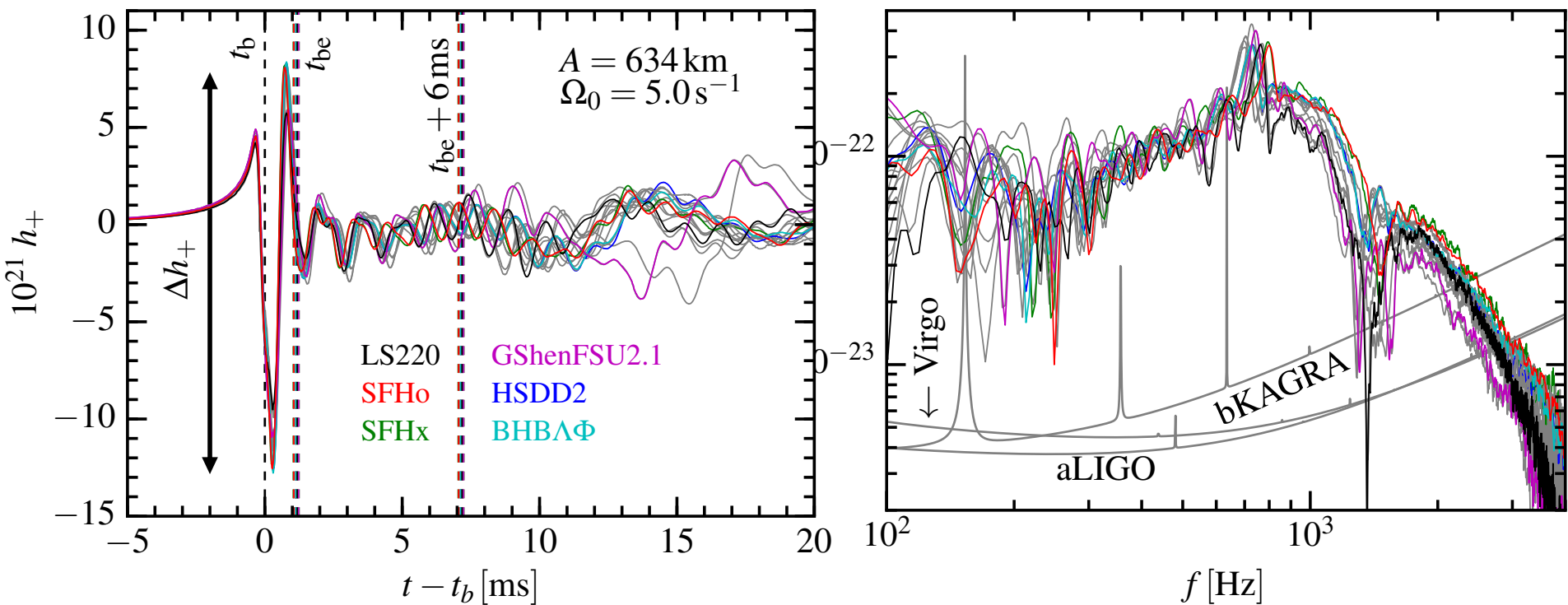




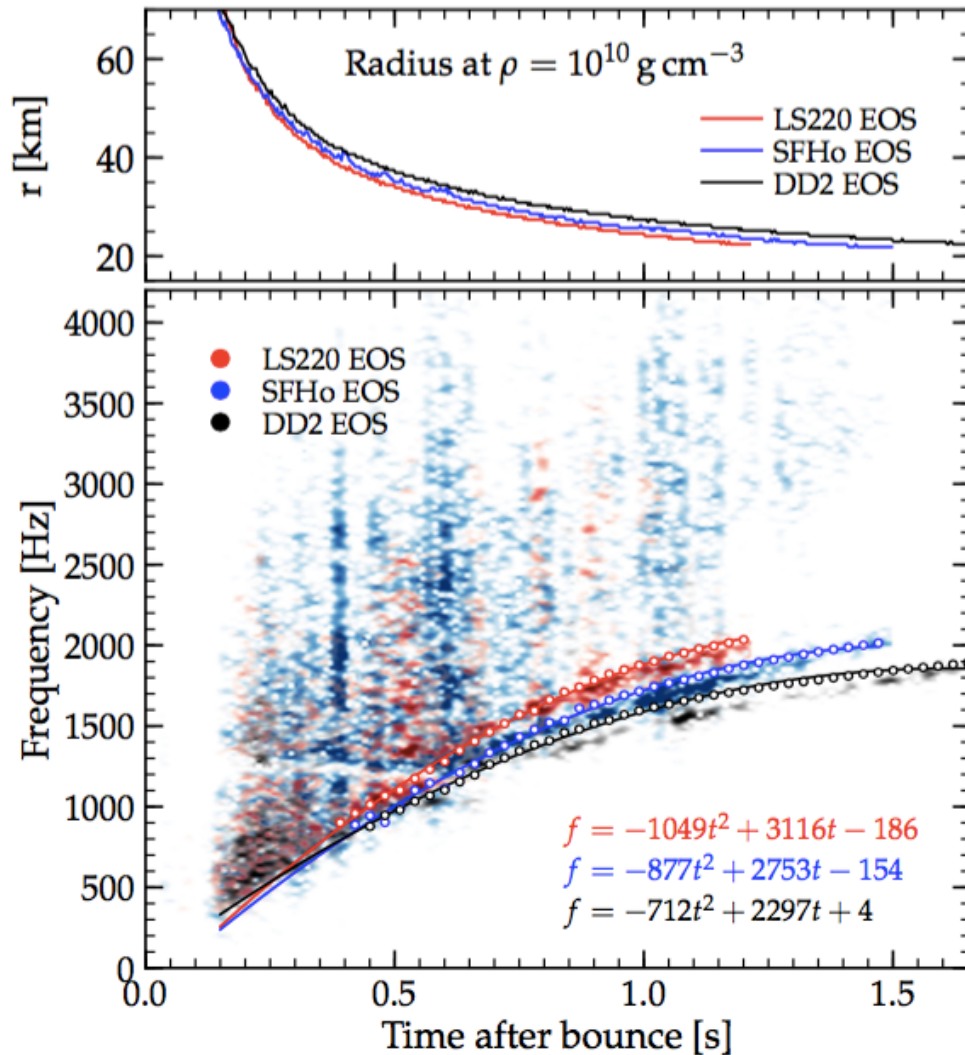
Weaker convection leads to weaker signal



EOS? Early Phase:

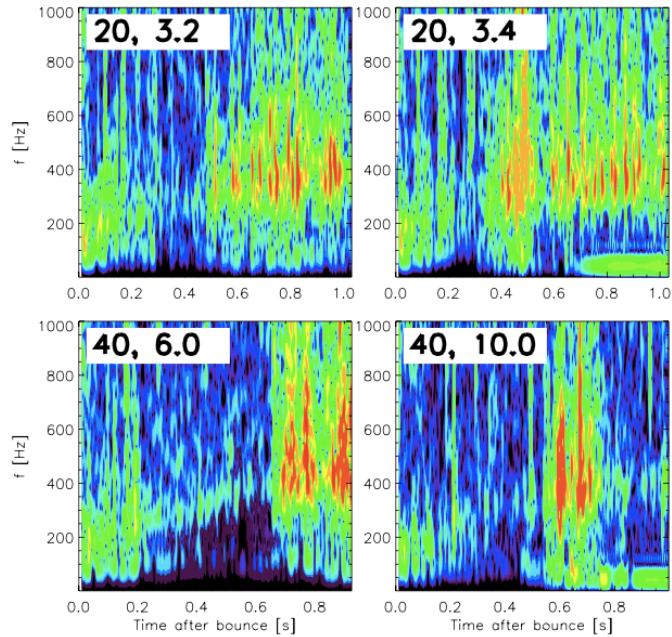


Dependence of the dominant GW frequency on the EOS

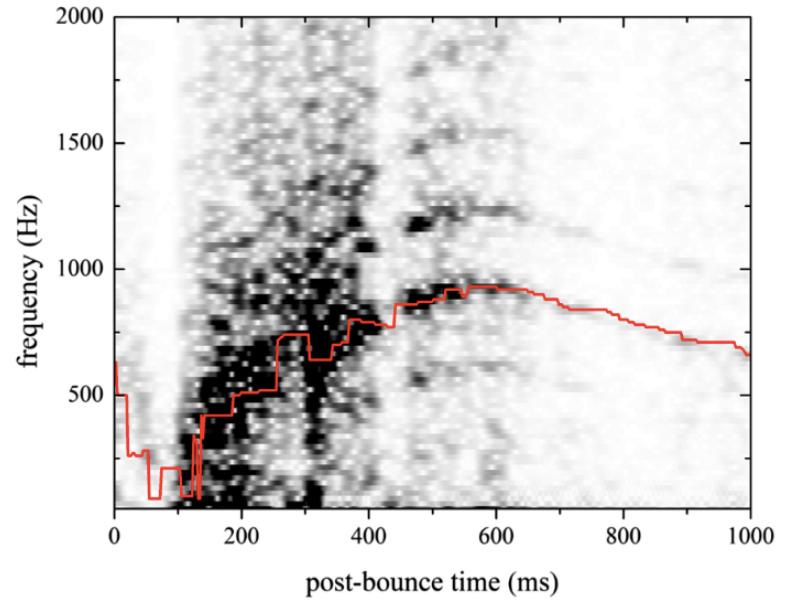


- Reflects the evolution of the PNS radius
- Captured reasonably well by the analysis
- Can be described by a quadratic function

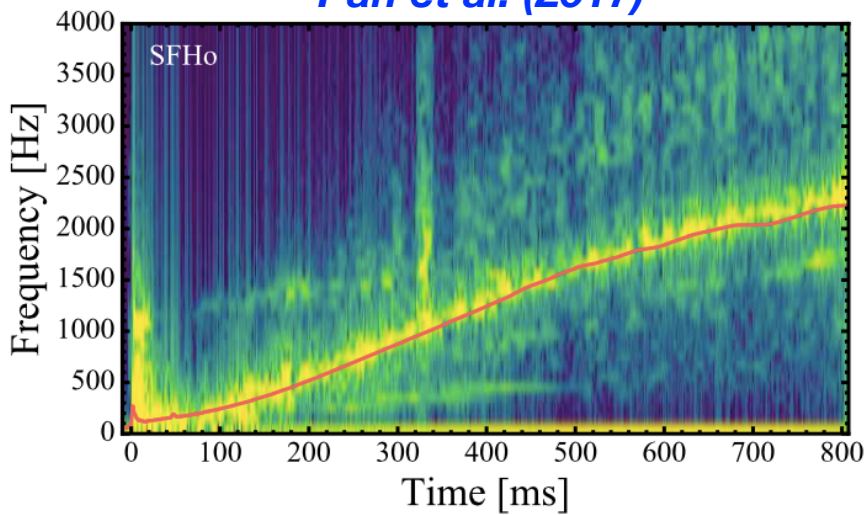
Murphy et al. (2009)



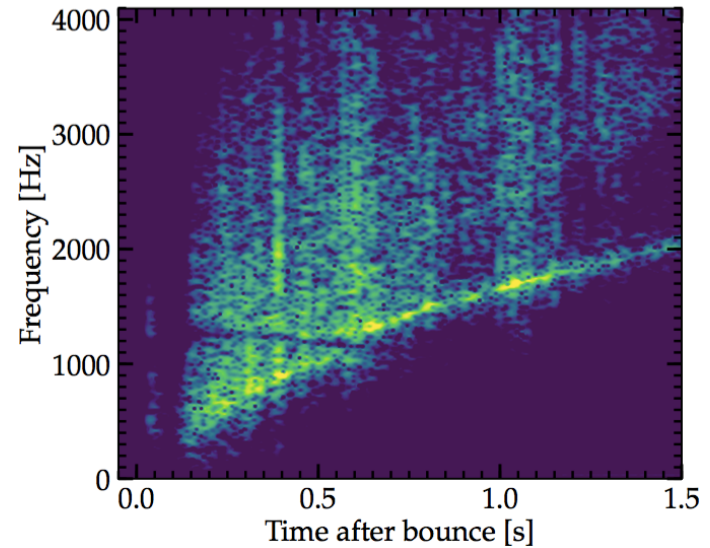
Yakunin et al. (2015)

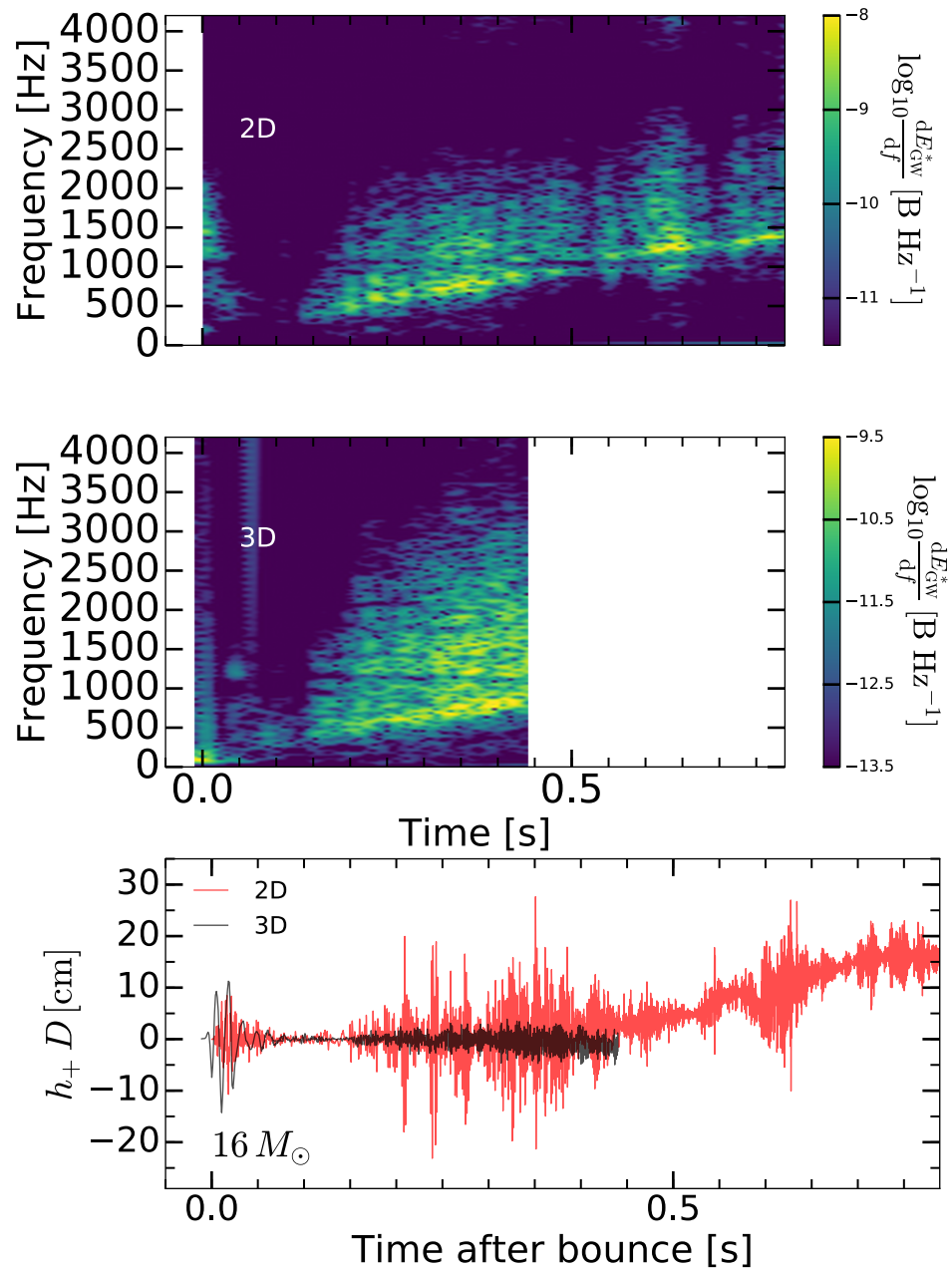


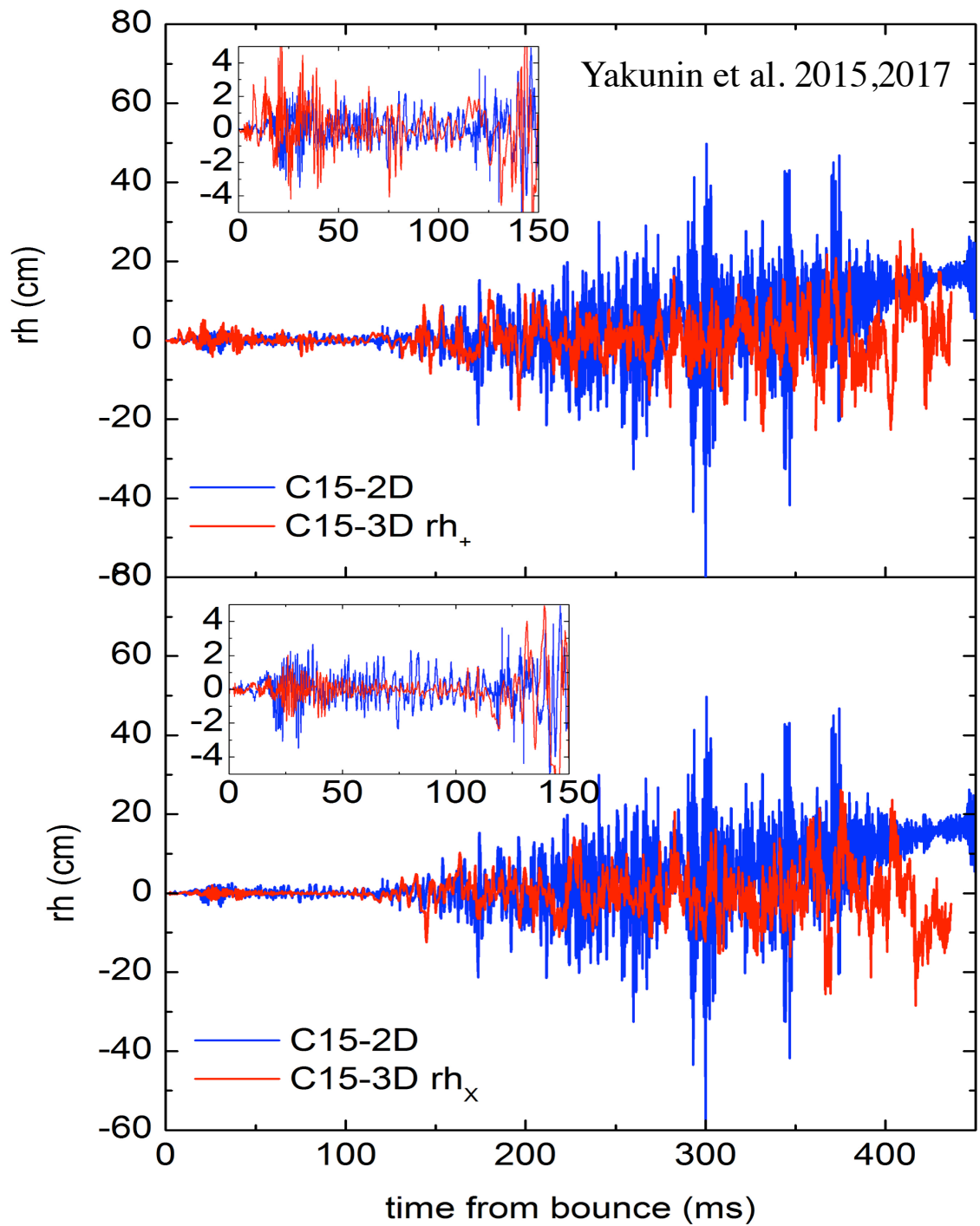
Pan et al. (2017)



Morozova et al. (2018)







Core-Collapse Theory: A Status Summary

- ♦ “Mazurek’s Law” Violated in Multi-D near Criticality
- ♦ Feedbacks severe in 1D, but in 2D/3D there is a strong dependence on microphysics and computational details
- ♦ Proximity to critical explosion curve amplifies effects of sub-dominant processes, etc.
- ♦ Can explain current differences between groups (!?)

- ♦ Multi-D is Key Enabler of explosion for (almost) all viable mechanisms
- ♦ Neutrino-driven convection > SASI (when object explodes to yield SN)
- ♦ SASI is not a mechanism - can’t generate much entropy
- ♦ “Ray-by-ray” is problematic in 2D
- ♦ 3D different from 2D (turbulent pressure, spectrum; scales)!
- ♦ GR important? Newtonian Models can explode

- ♦ Various heating processes (in-medium/many-body, inelastic on electrons, inelastic on nucleons) add “non-linearly”
- ♦ Structure factor/many-body corrections! Neutrino-matter interactions!

- ♦ Progenitor profiles/structure important! (e.g., Meakin & Arnett; Couch et al. 2015; B. Muller et al. 2016); Seed Perturbations, Density profiles, Si/O shelves?
- ♦ Rotation!?
- ♦ Crucial role for microphysics - many-body/structure-factor corrections, inelastic scattering; when near critical curve, small effects are amplified - (partial) origin of differences between groups